

NPAC course on Astroparticles

CONCLUSIONS

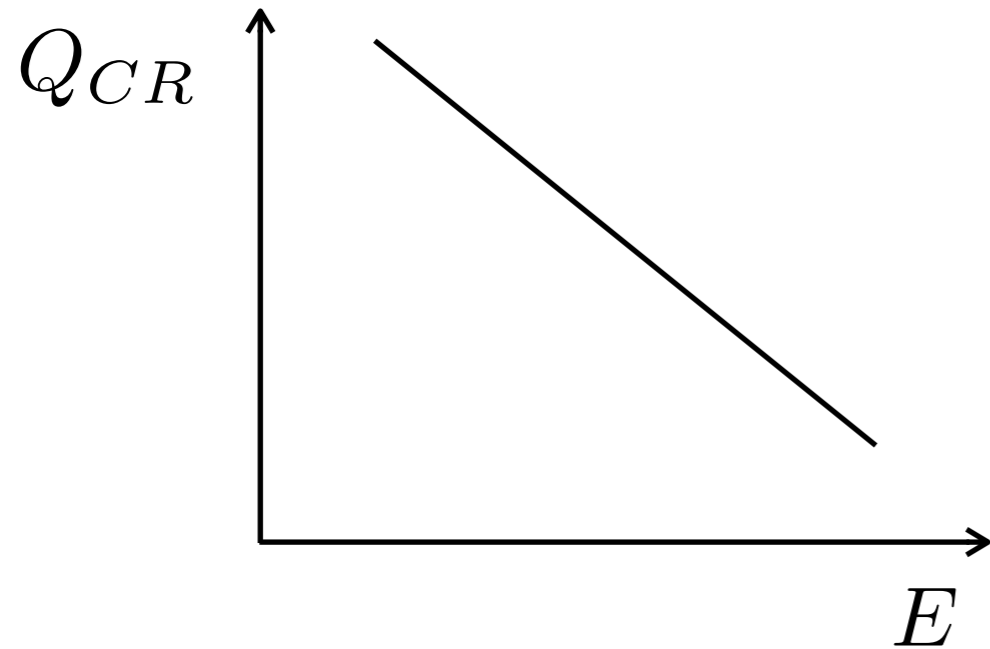
NPAC course on Astroparticles

~~CONCLUSIONS~~

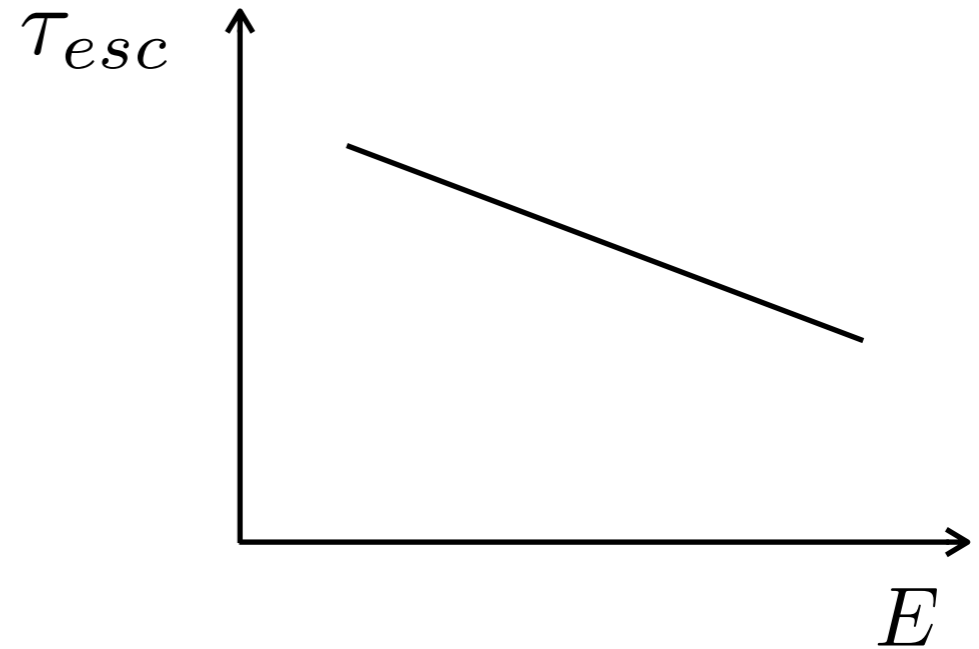
(incomplete list of) PROBLEMS...

The two-power-laws scenario

$$Q_{CR}(E) \propto E^{-\delta}$$

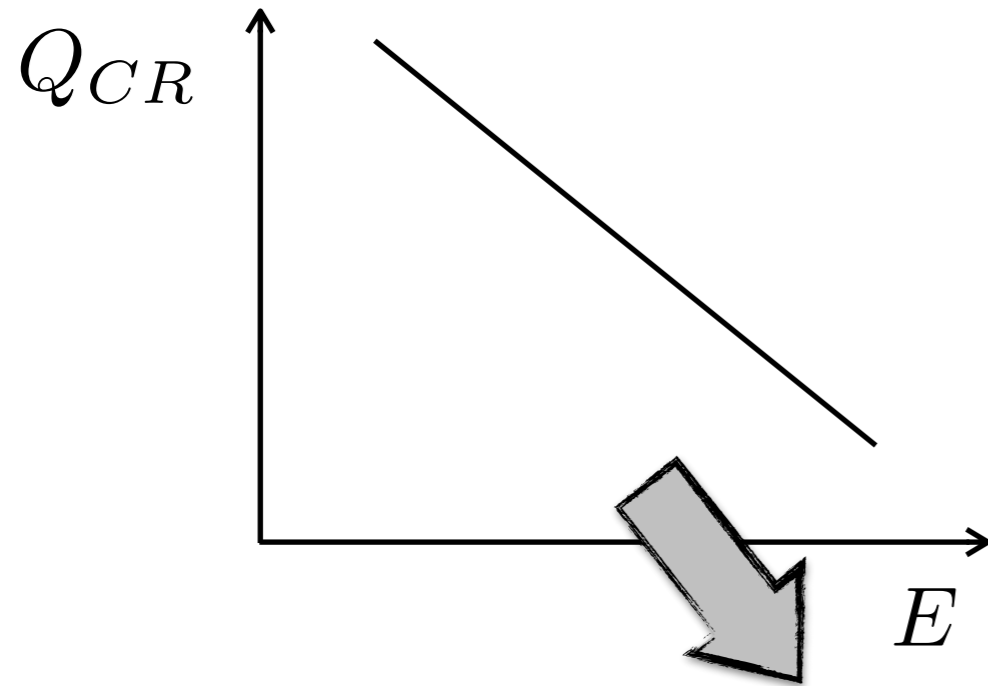


$$\tau_{esc}(E) \propto E^{-\alpha}$$

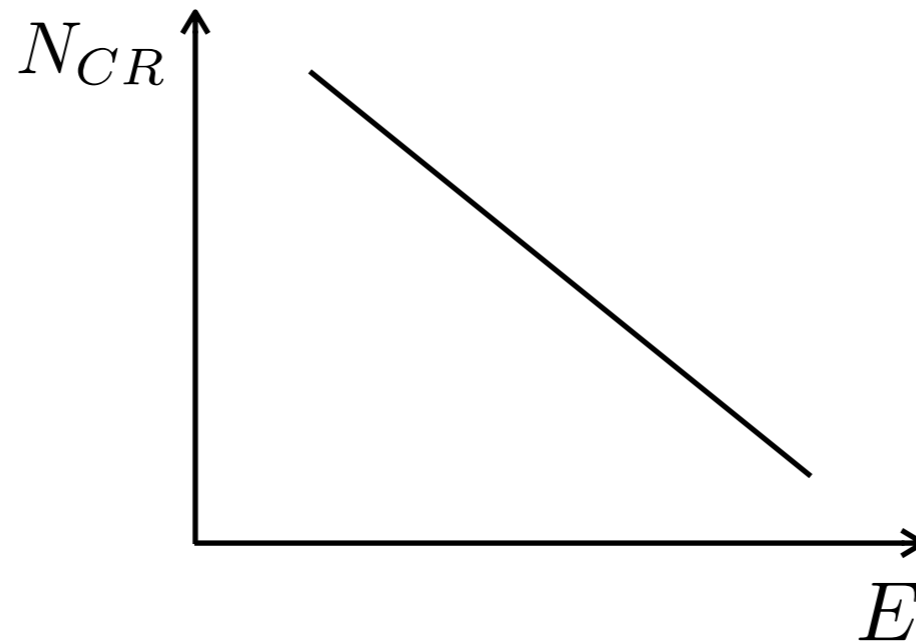
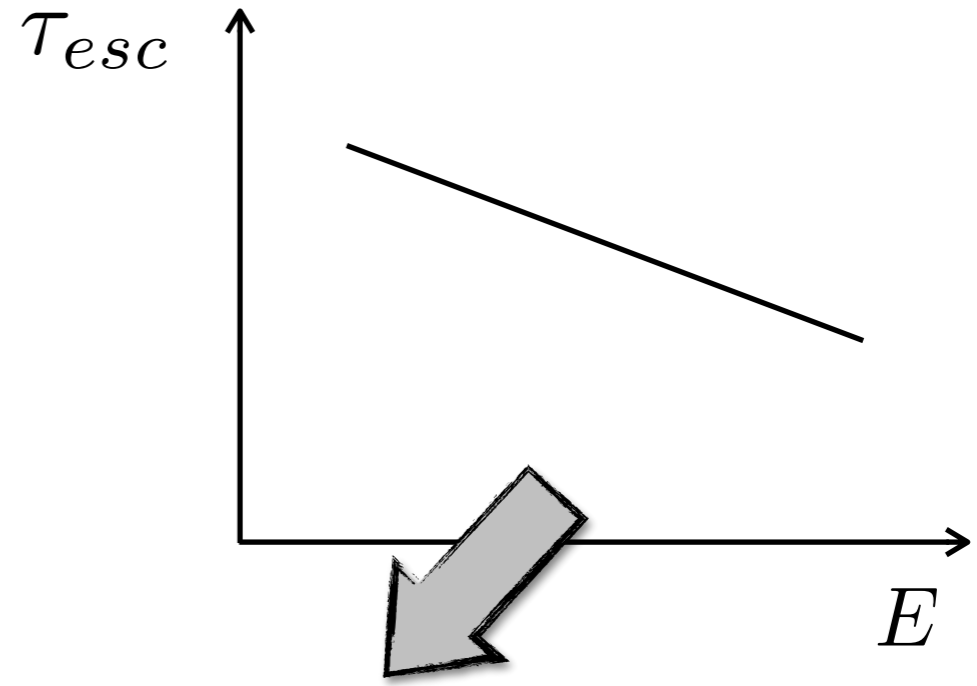


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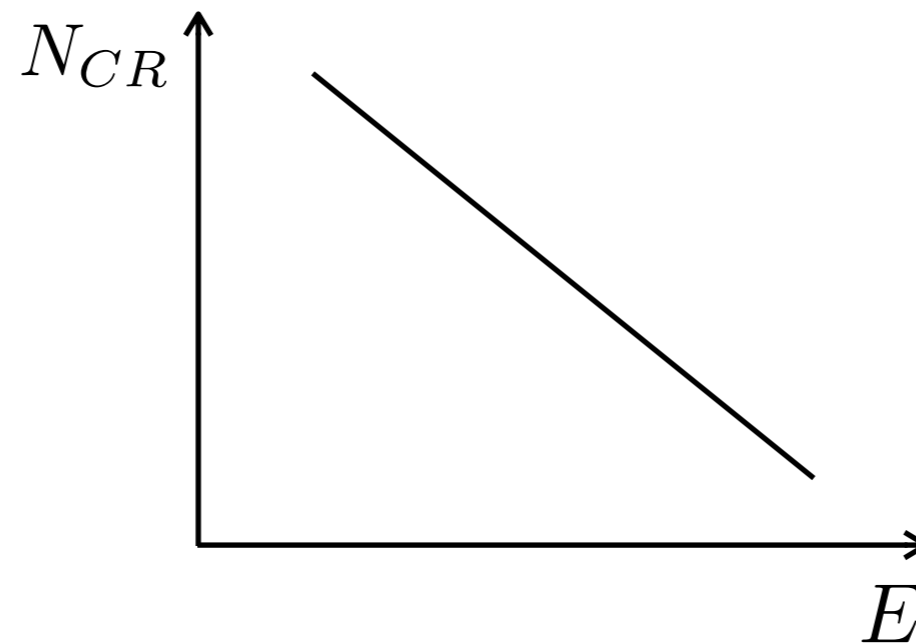
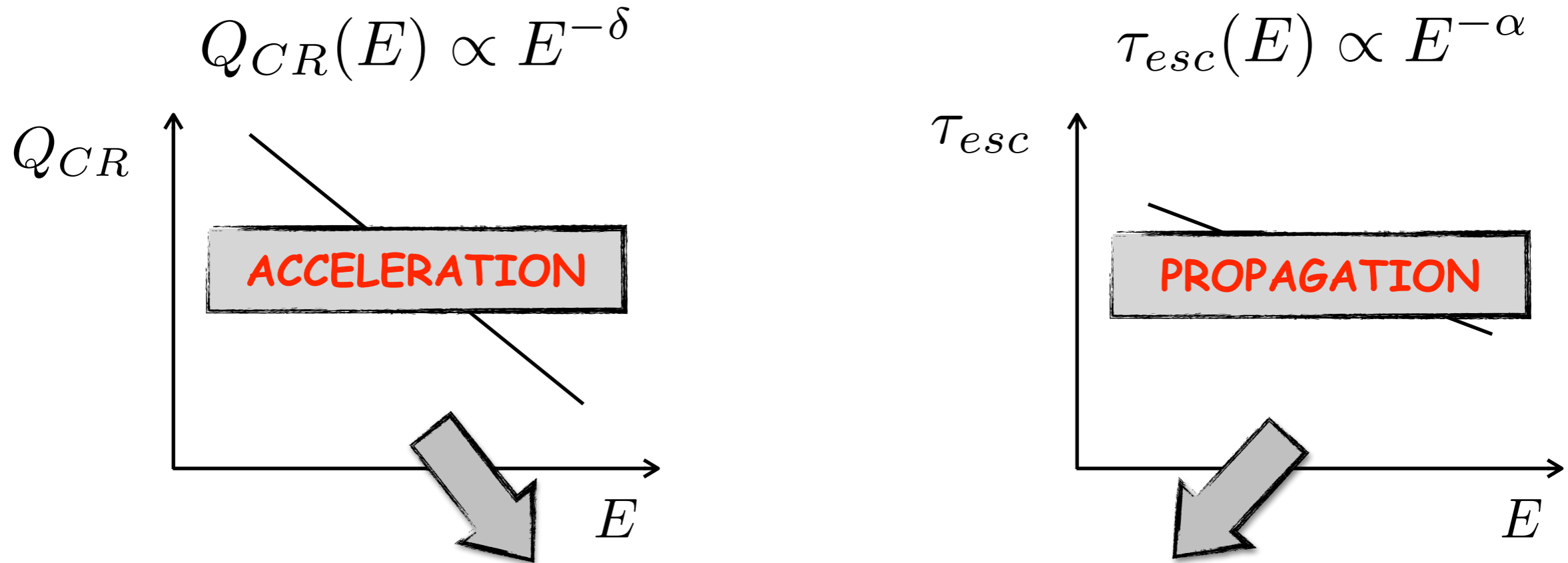


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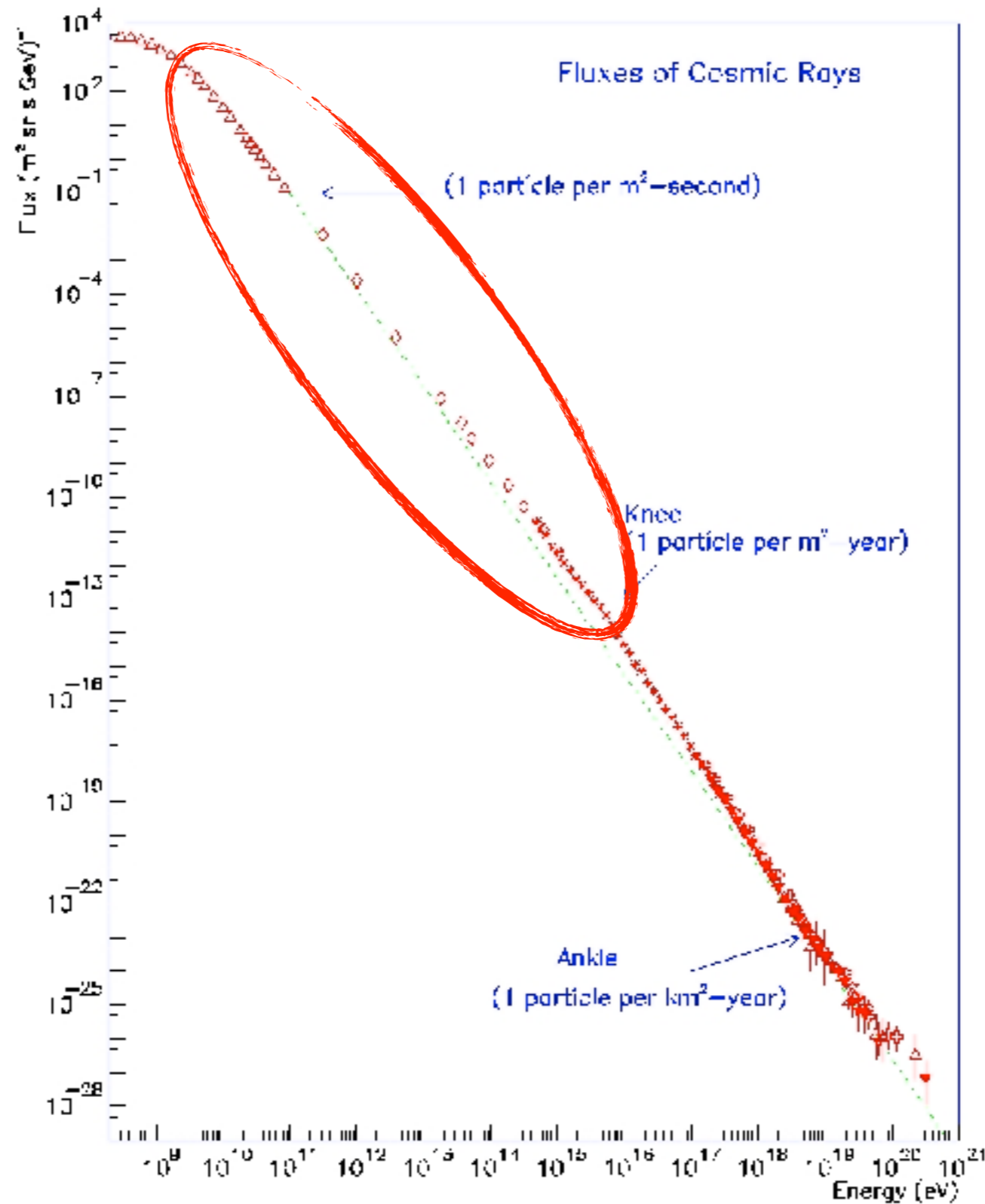
$$N_{CR}(E) = Q_{CR}(E) \times \tau(E) \propto E^{-\delta-\alpha}$$

The two-power-laws scenario



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The observed cosmic ray spectrum



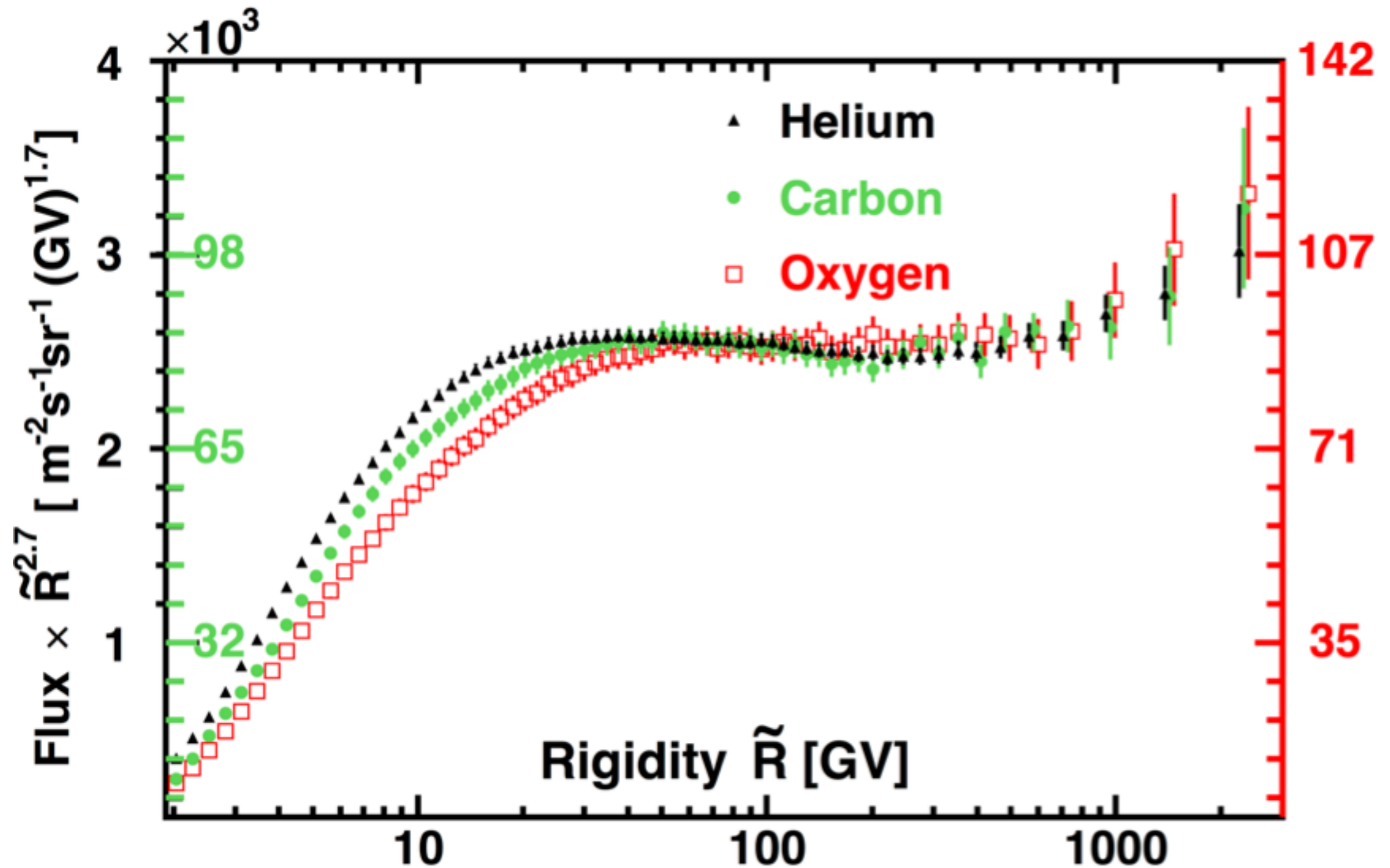
Real life is much harder than that..



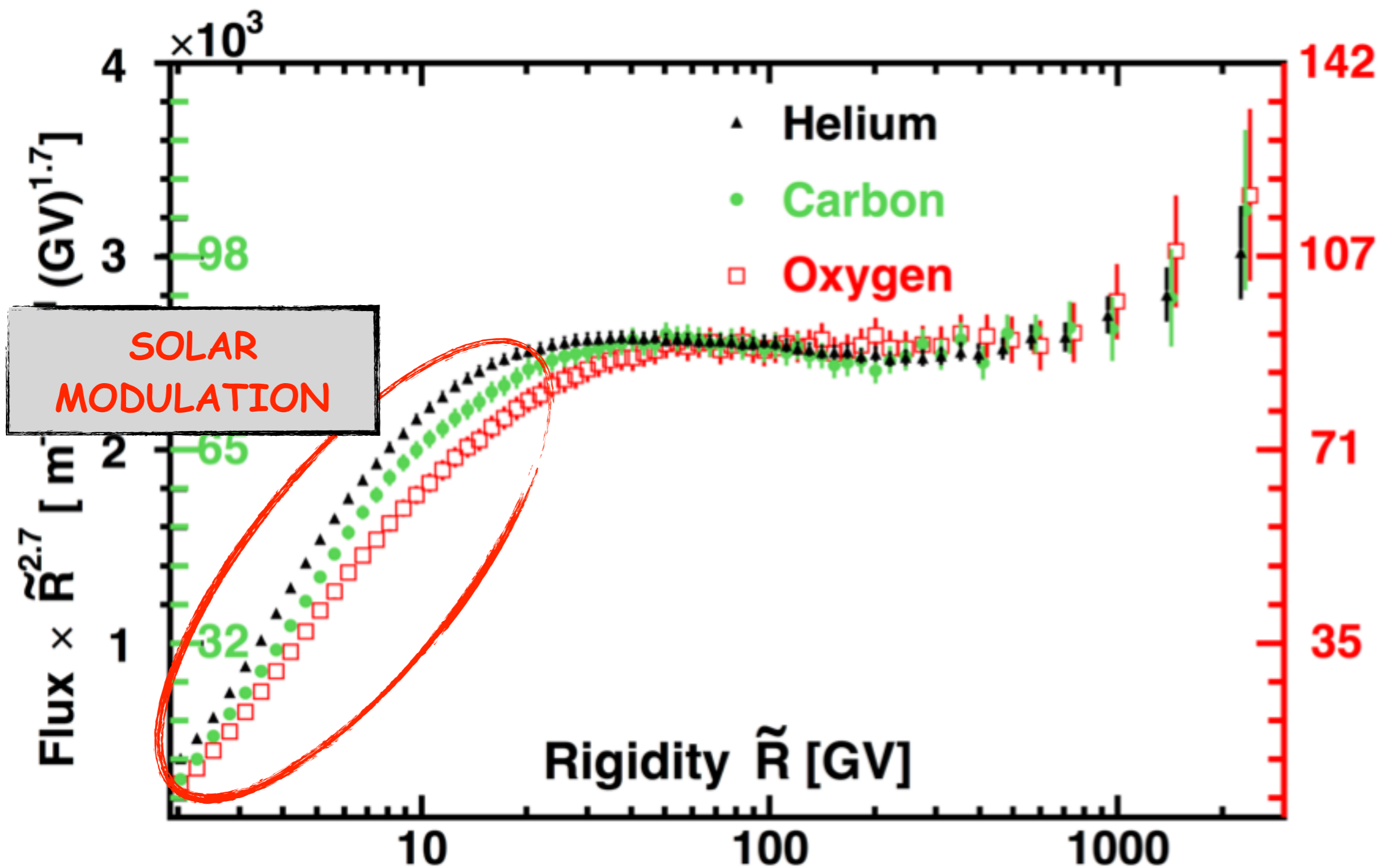
PAMELA and AMS-02



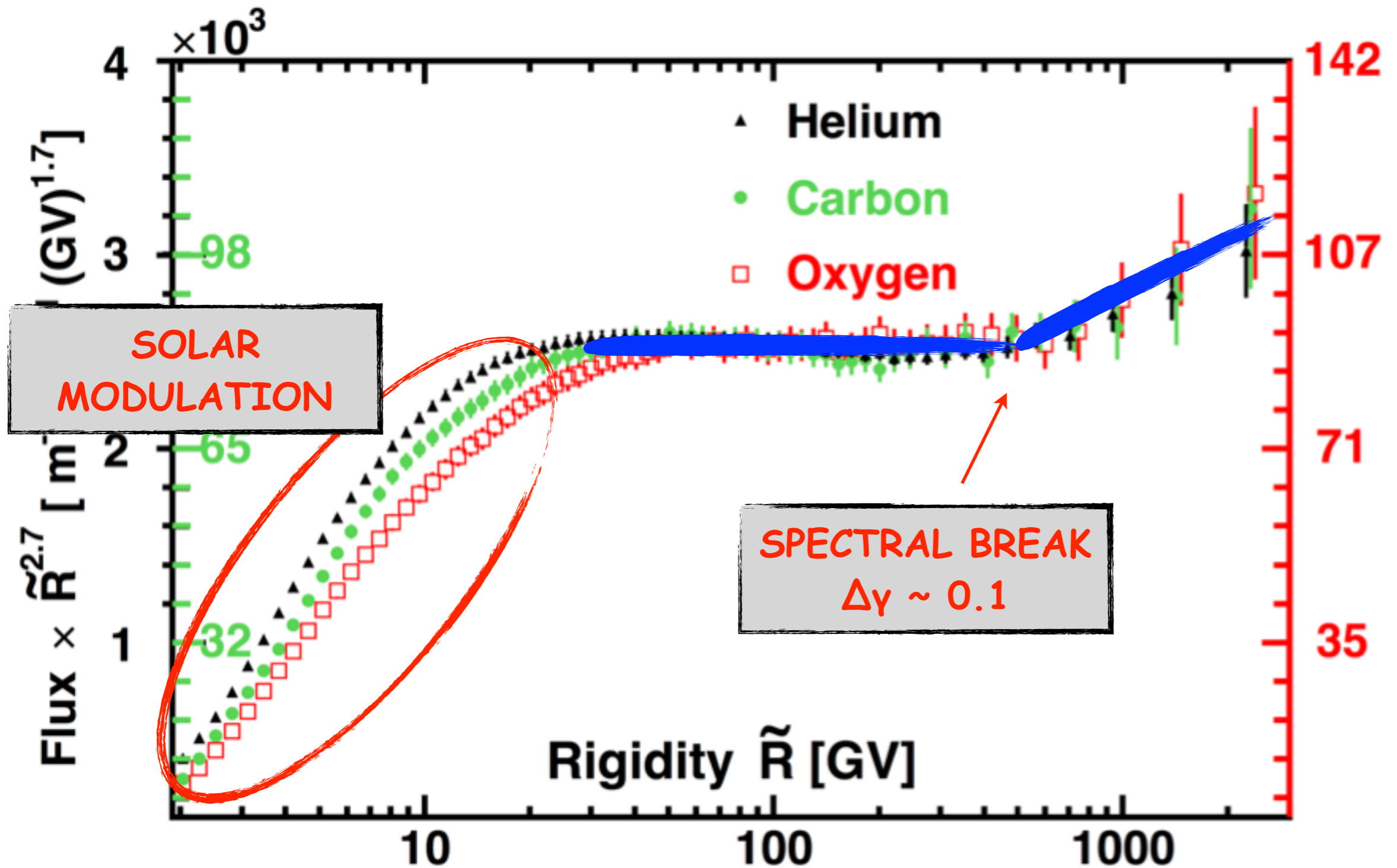
Spectral breaks at 300 GV



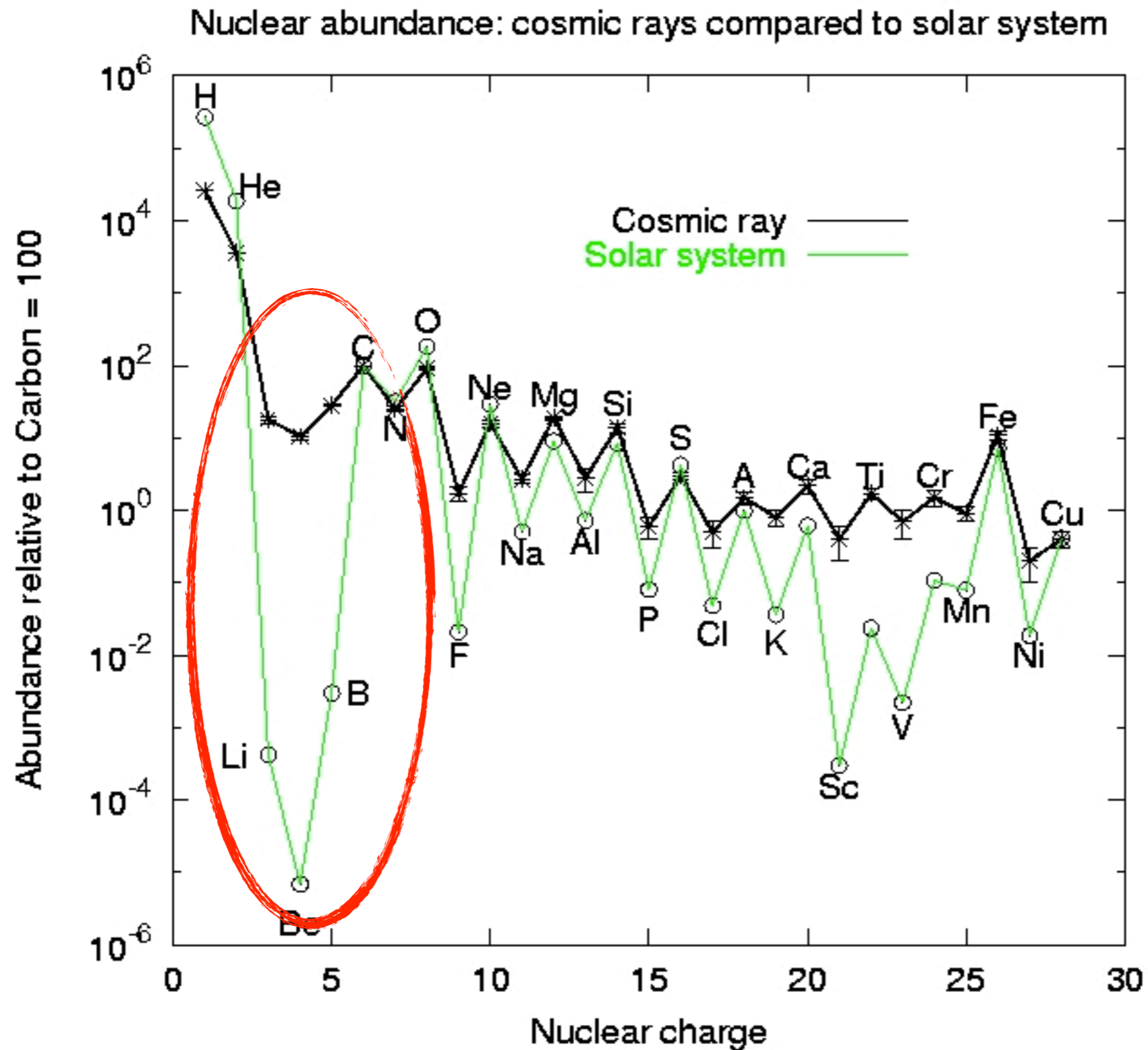
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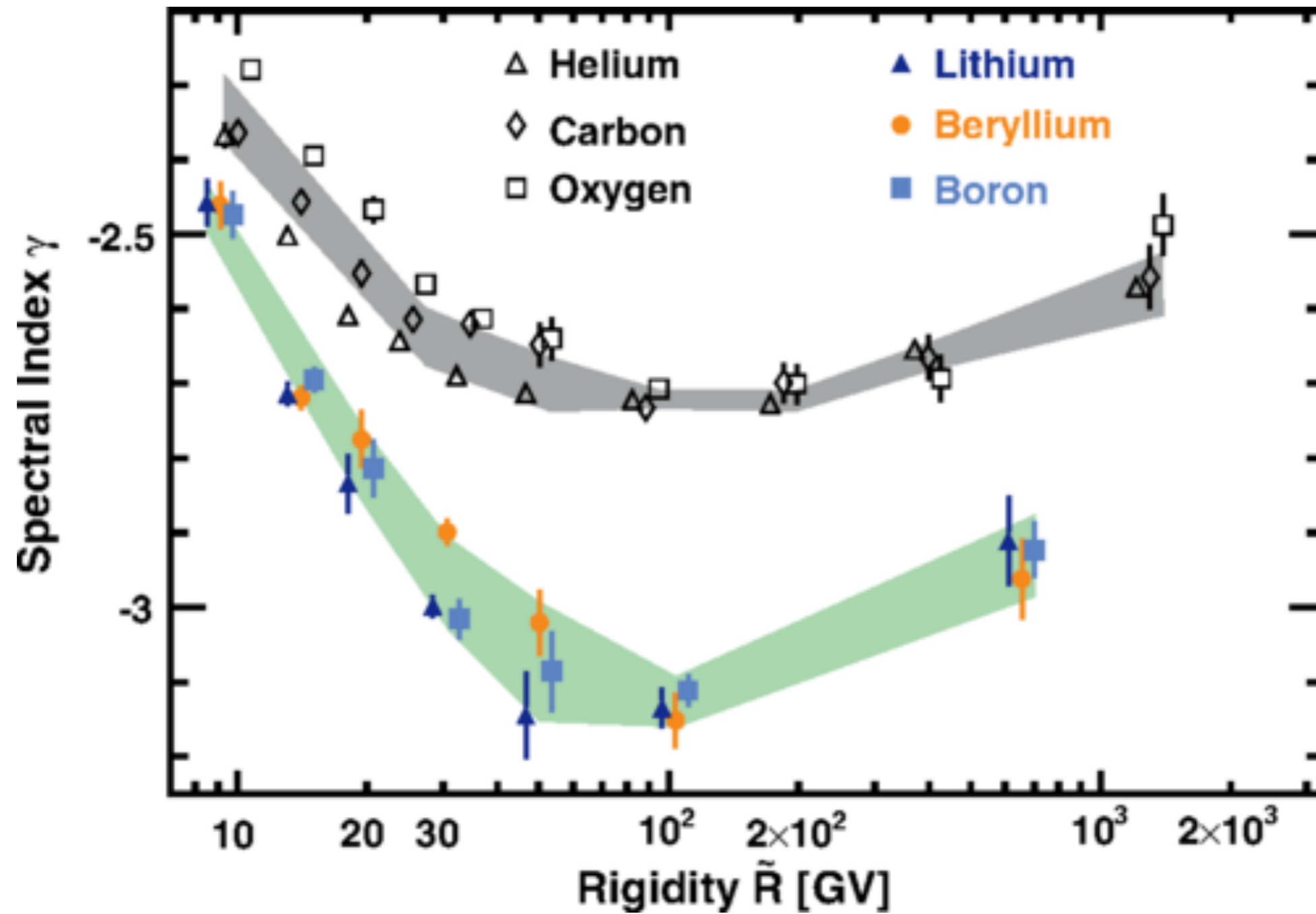
Spectral breaks at 300 GV



Acceleration or propagation?

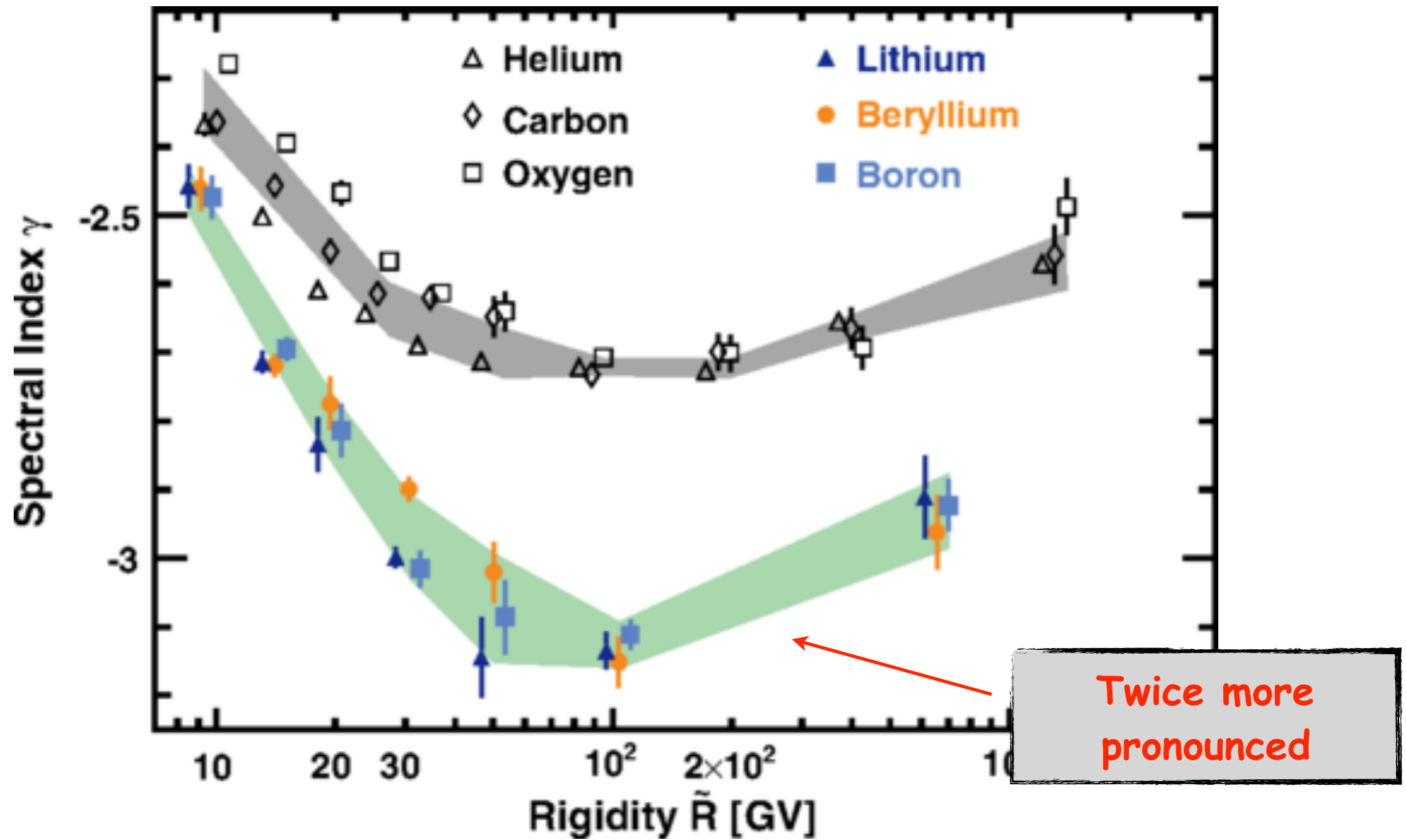


Propagation!



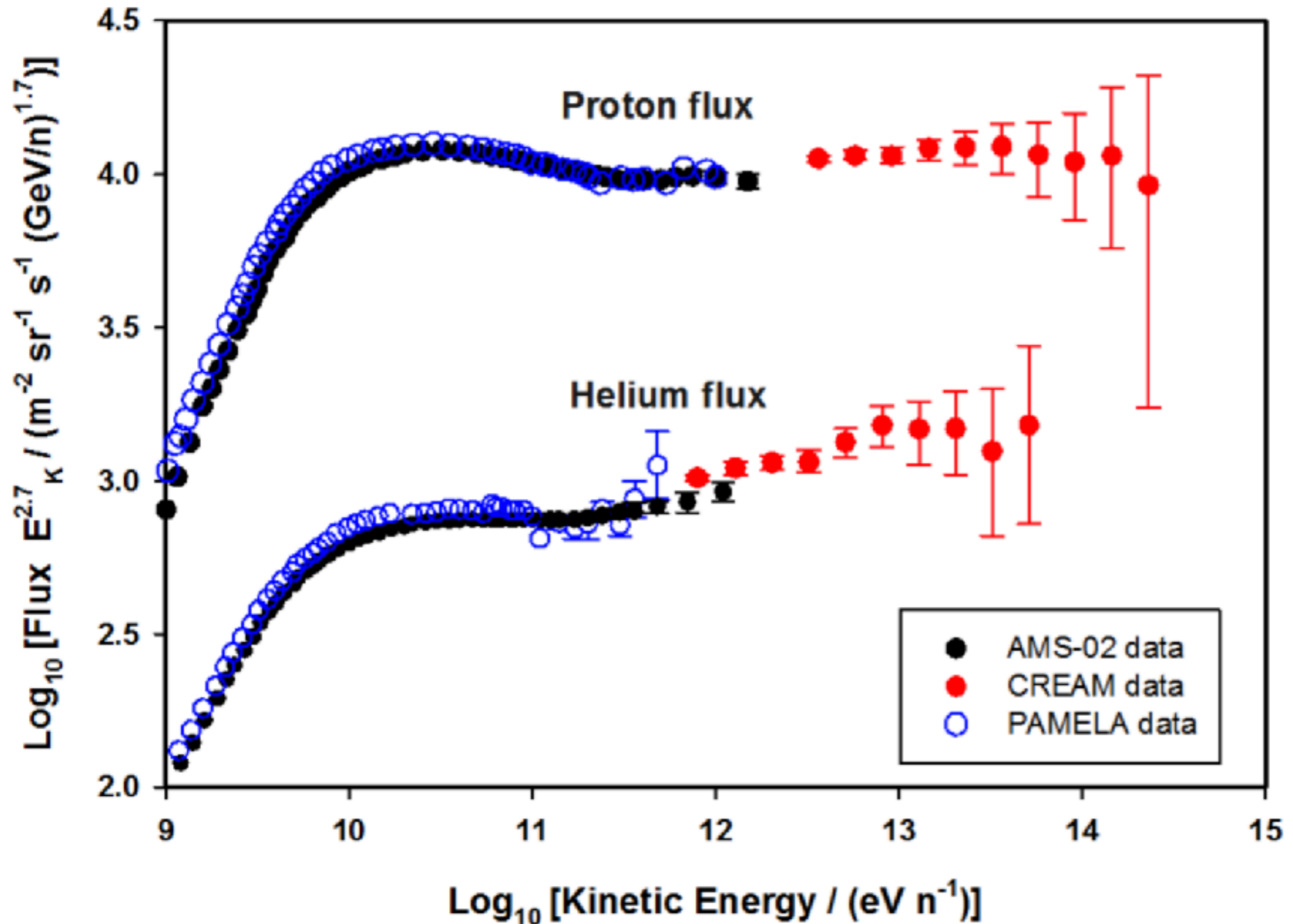
What happens at 300 GV?

Propagation!

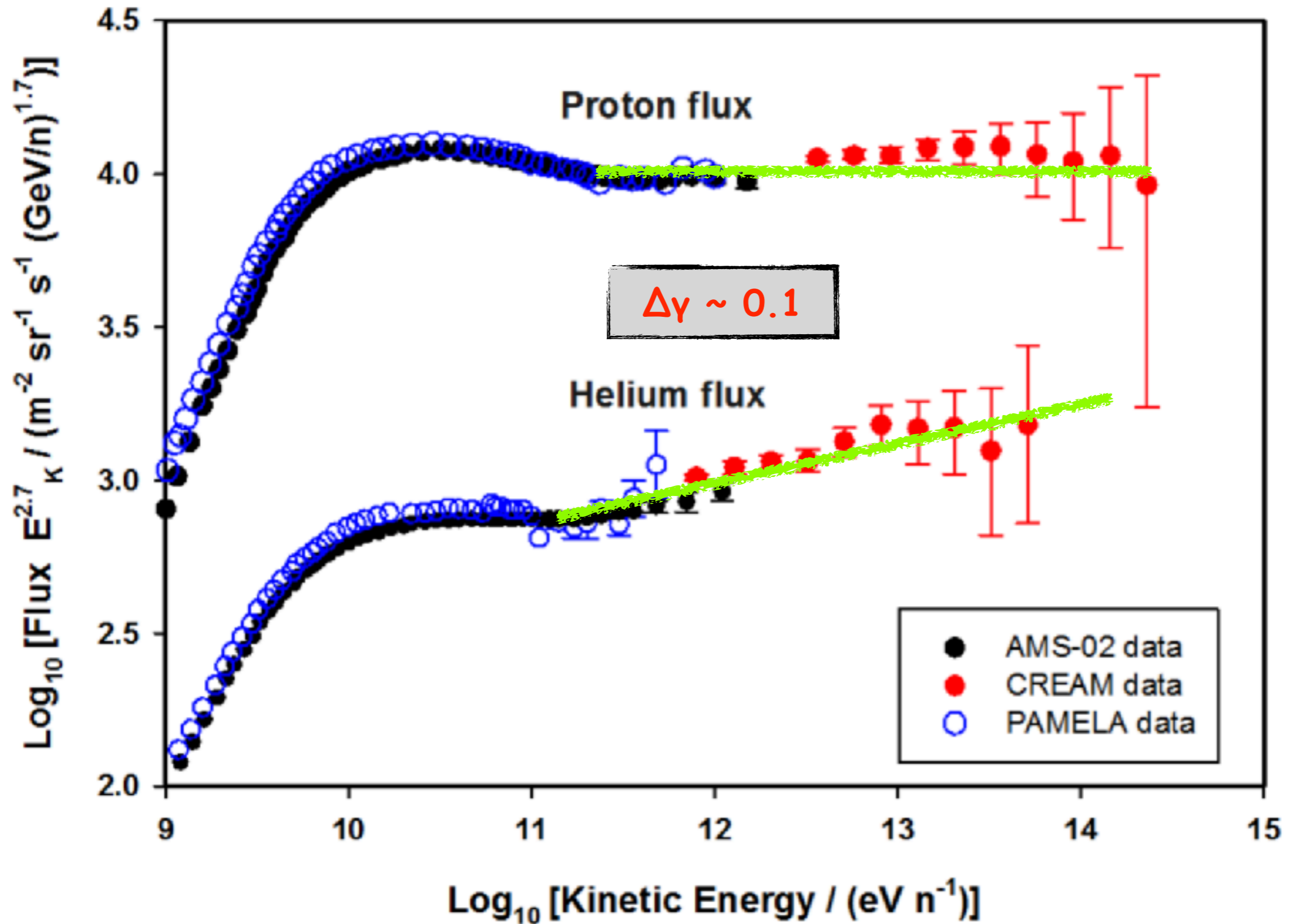


What happens at 300 GV?

H spectrum is softer than He... (?!?!)

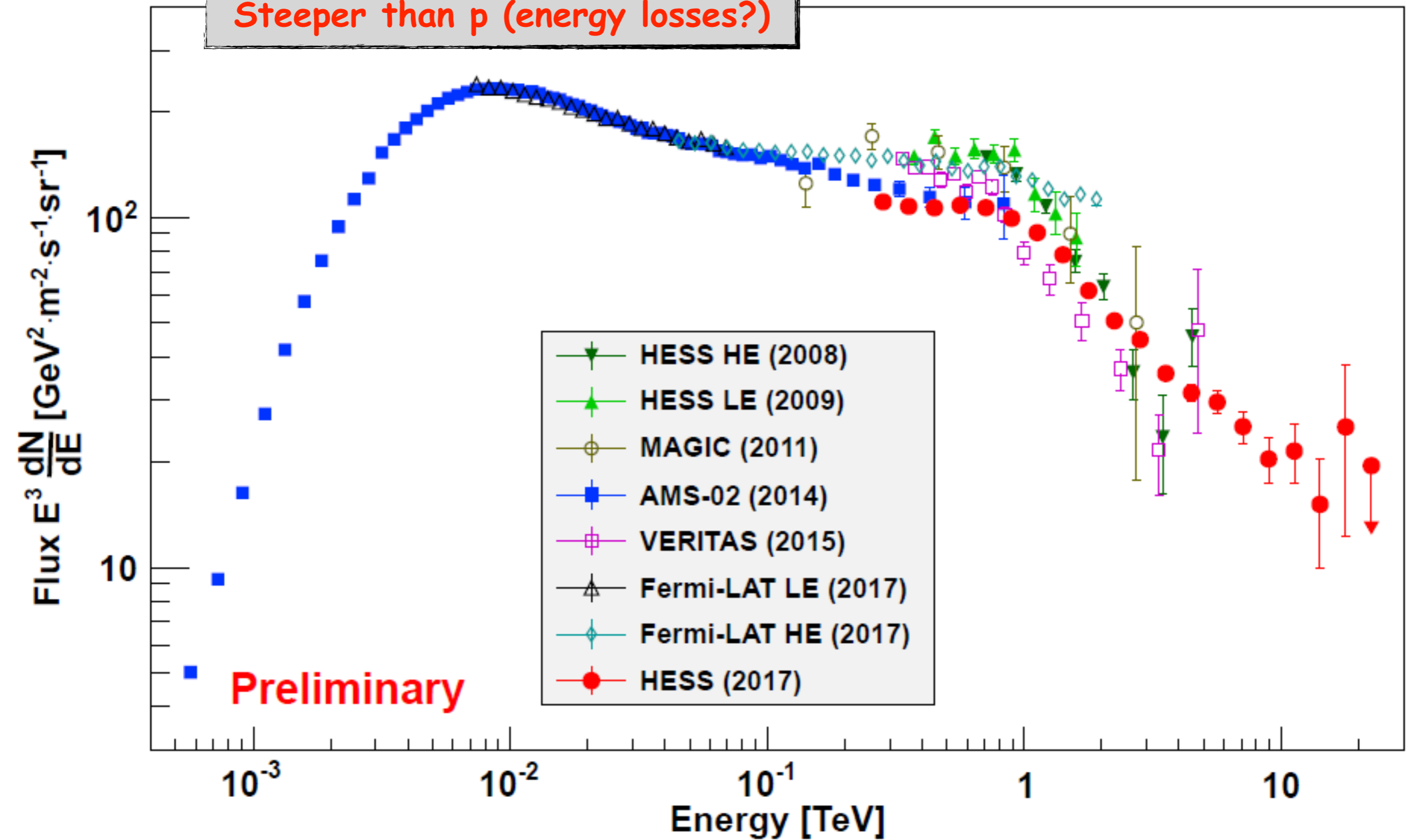


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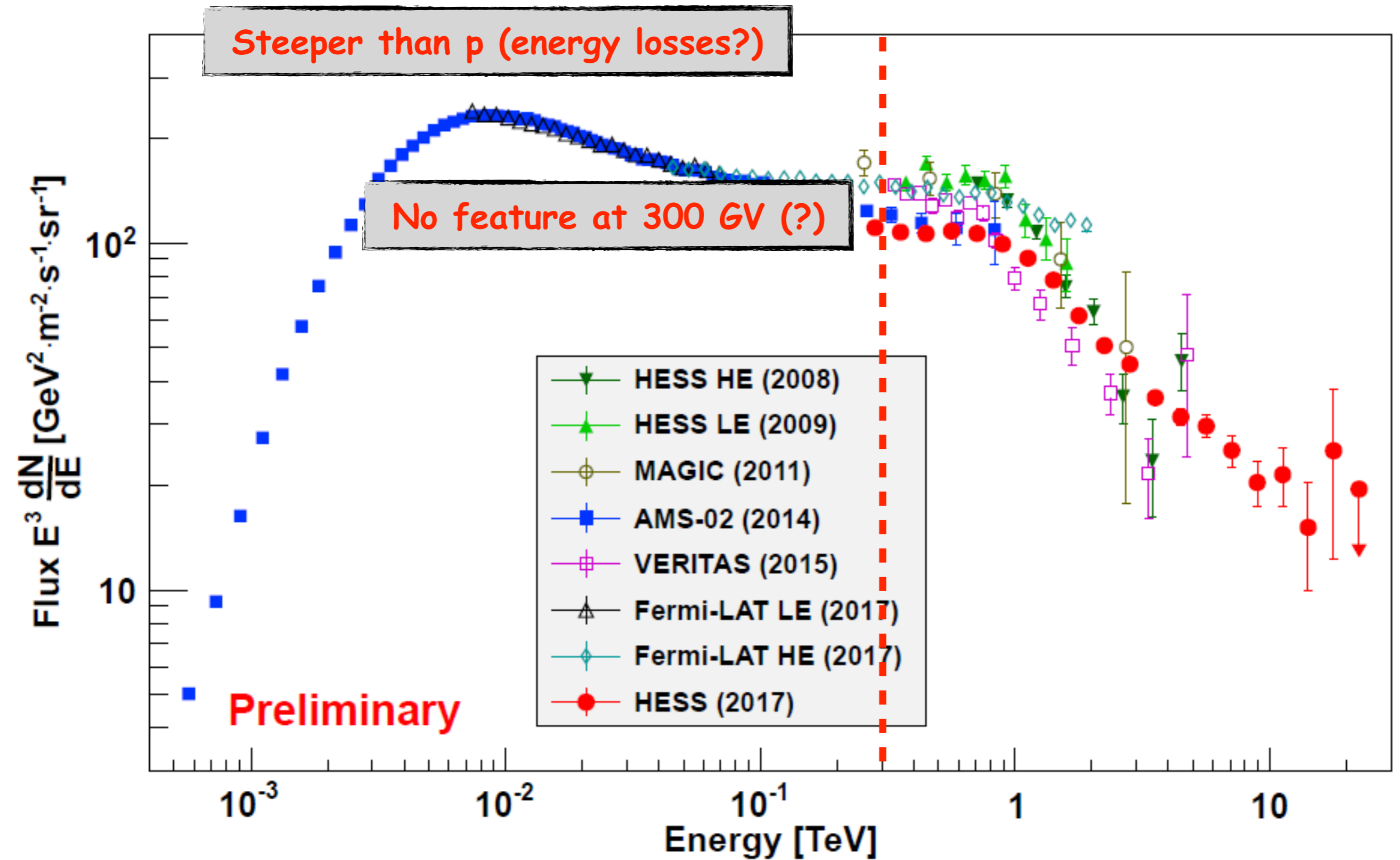


Electron spectrum

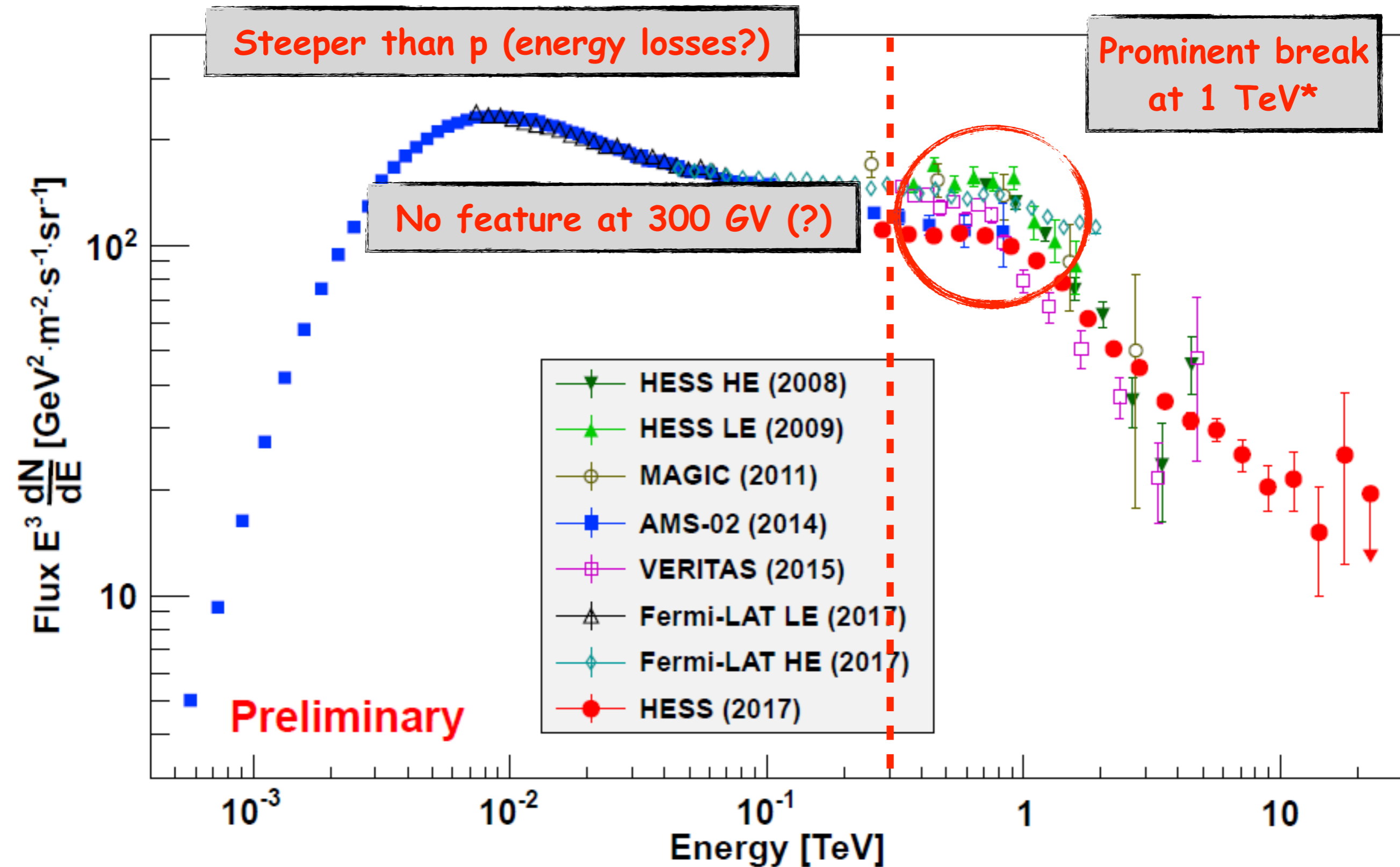
Steeper than p (energy losses?)



Electron spectrum

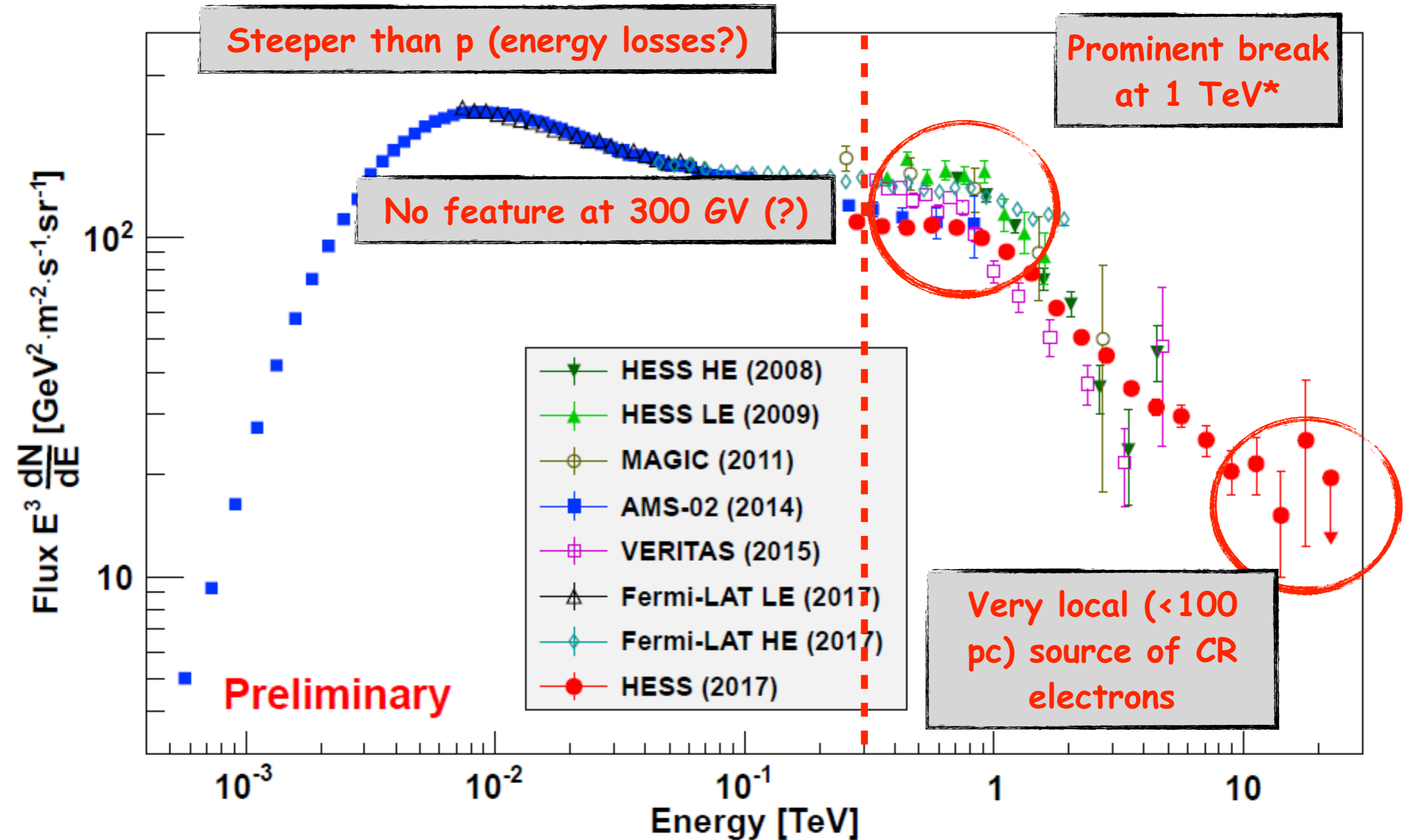


Electron spectrum



* most pronounced spectral feature in the entire spectrum of cosmic ray particles!

Electron spectrum



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Positrons

Neutrinos/antineutrinos & electrons/positrons are also produced in pp interactions

$$p + p \rightarrow p + p + \pi^0 + \pi^+ + \pi^-$$

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$$\left\{ \begin{array}{l} \pi^\pm \rightarrow \mu^\pm + \nu_\mu(\bar{\nu}_\mu) \\ \mu^\pm \rightarrow e^\pm + \bar{\nu}_\mu(\nu_\mu) + \nu_e(\bar{\nu}_e) \end{array} \right.$$

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Final products of proton-proton interactions are not only **gamma ray photons** but also **neutrinos, anti-neutrinos, electrons** and **positrons**

$$E_e \approx E_\nu \approx \frac{E_p}{20}$$

Positrons

Neutrinos/antineutrinos & electrons/positrons are also produced in pp interactions

$$p + p \rightarrow p + p + \pi^0 + \pi^+ + \pi^-$$

neutral and charged pions produced with the same probability (1/3,1/3,1/3)

$$\pi^0 \rightarrow \gamma + \gamma$$

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CR protons/electrons
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$$N_p(E) = Q_p(E) \times \tau_{esc}(E) \propto E^{-\delta-\alpha}$$

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Positrons come from
p-p interactions ->

$$Q_{e^+}(E) \propto N_p(E)$$

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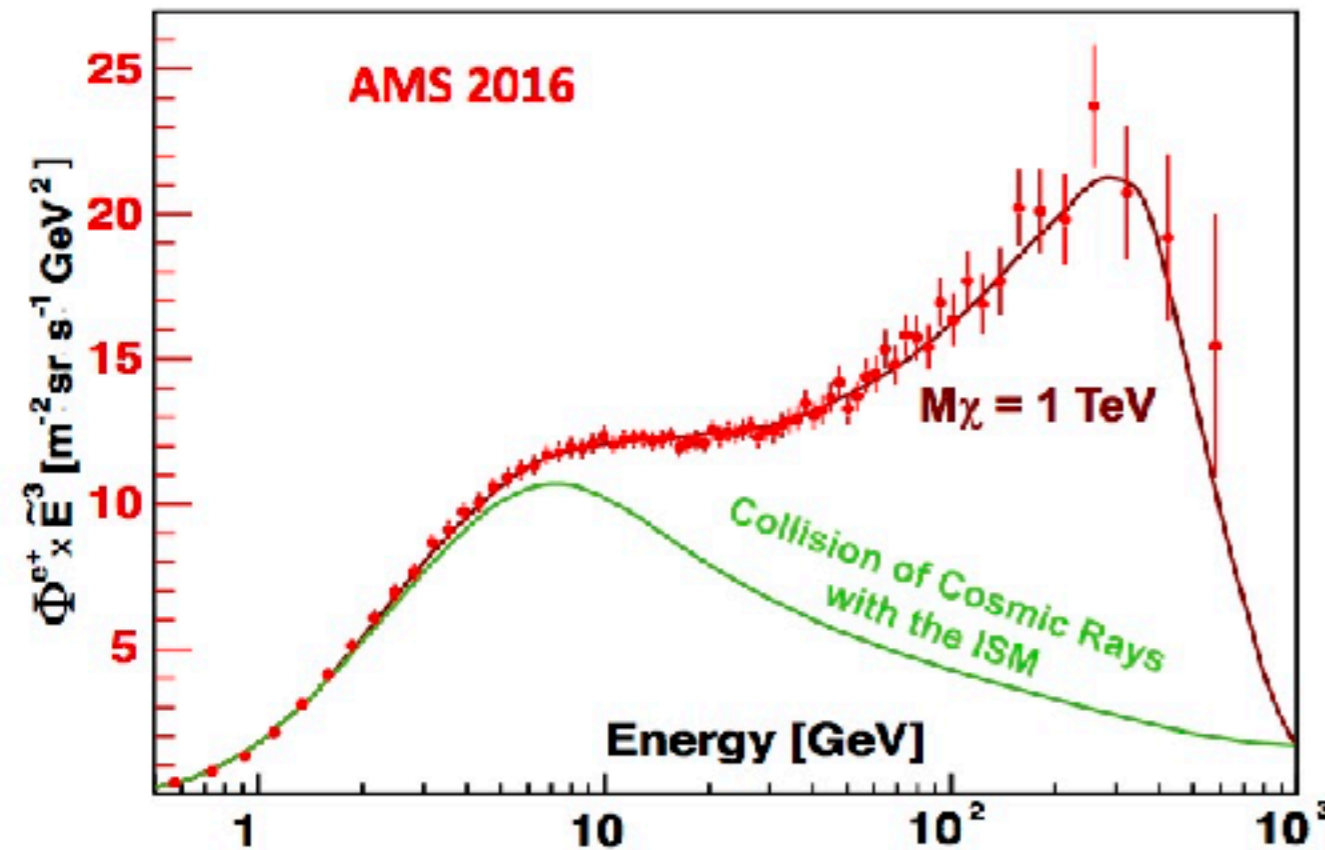
$$Q_{e^+}(E) \propto N_p(E)$$

Positrons and electrons behave in the same way ->

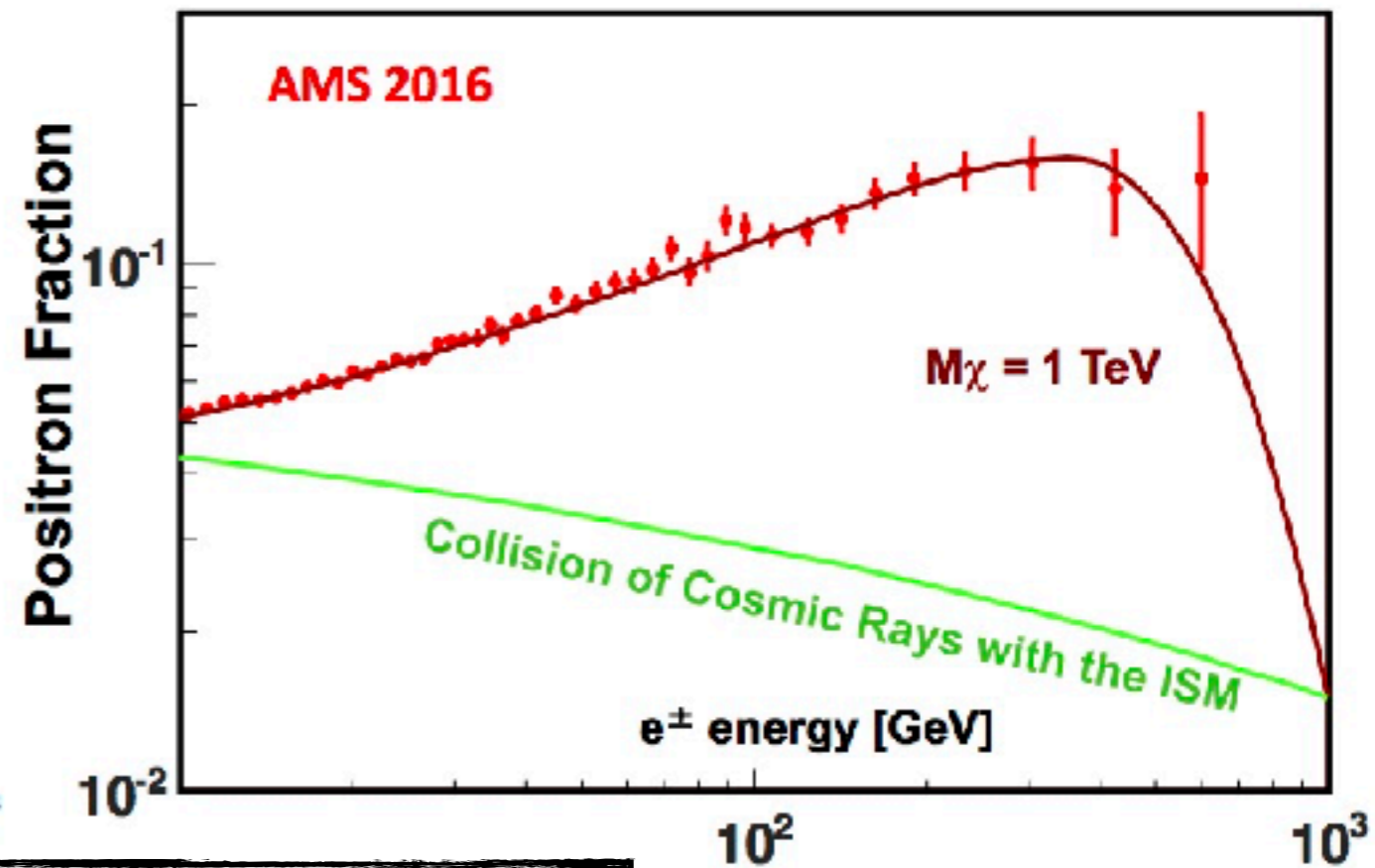
$$\frac{N_{e^+}(E)}{N_{e^-}(E)} = \frac{Q_{e^+}(E)}{Q_{e^-}(E)} \propto E^{-\alpha}$$

Positron fraction

Positron Spectrum



Positron Fraction



Nearby source of CR positrons? (pulsars?)

Dark matter decay?

Anisotropy

isotropic part

$$\int_{-1}^1 d\mu \mu v f^{(0)} = 0$$

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increases with E

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points towards
Galactic centre

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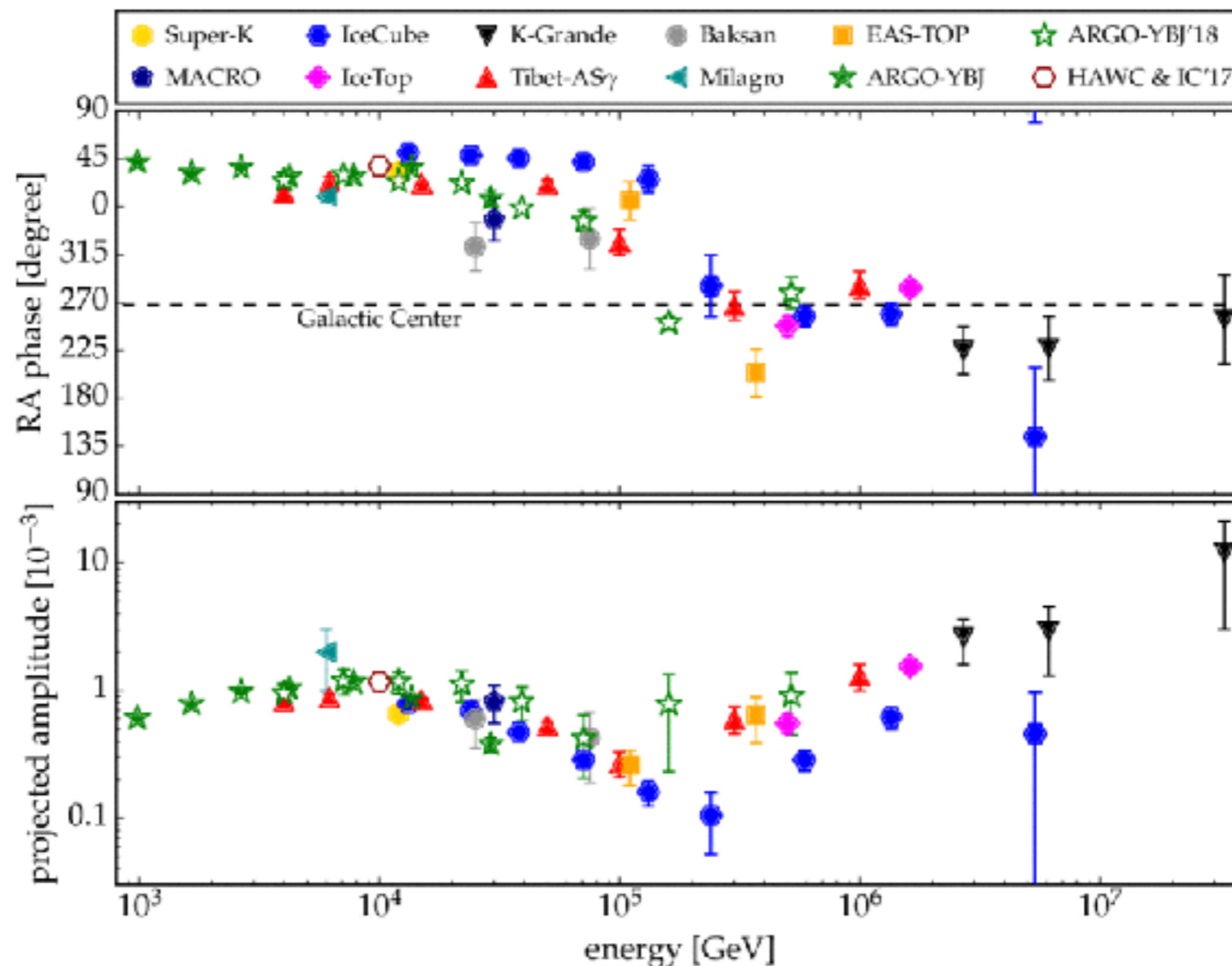
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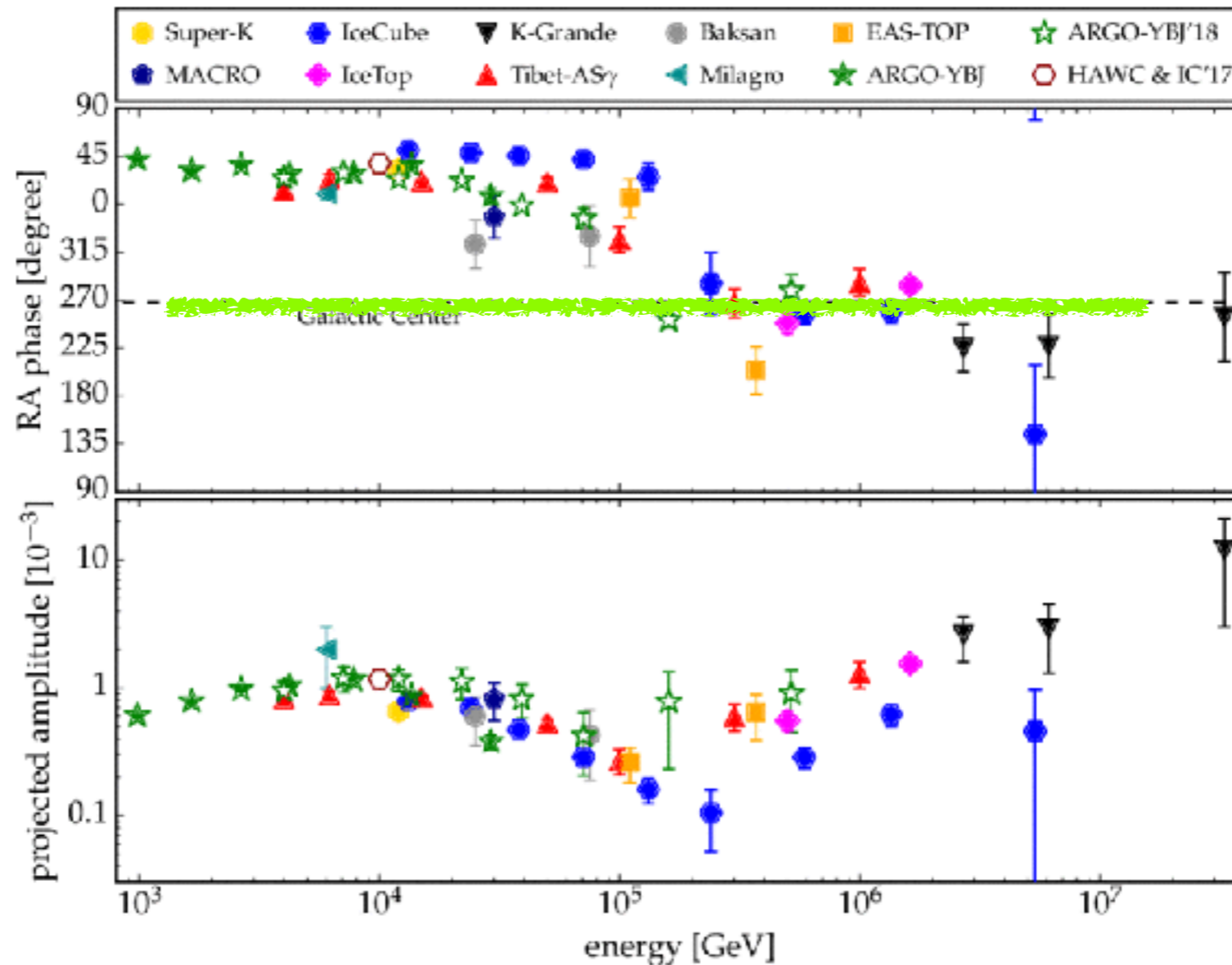
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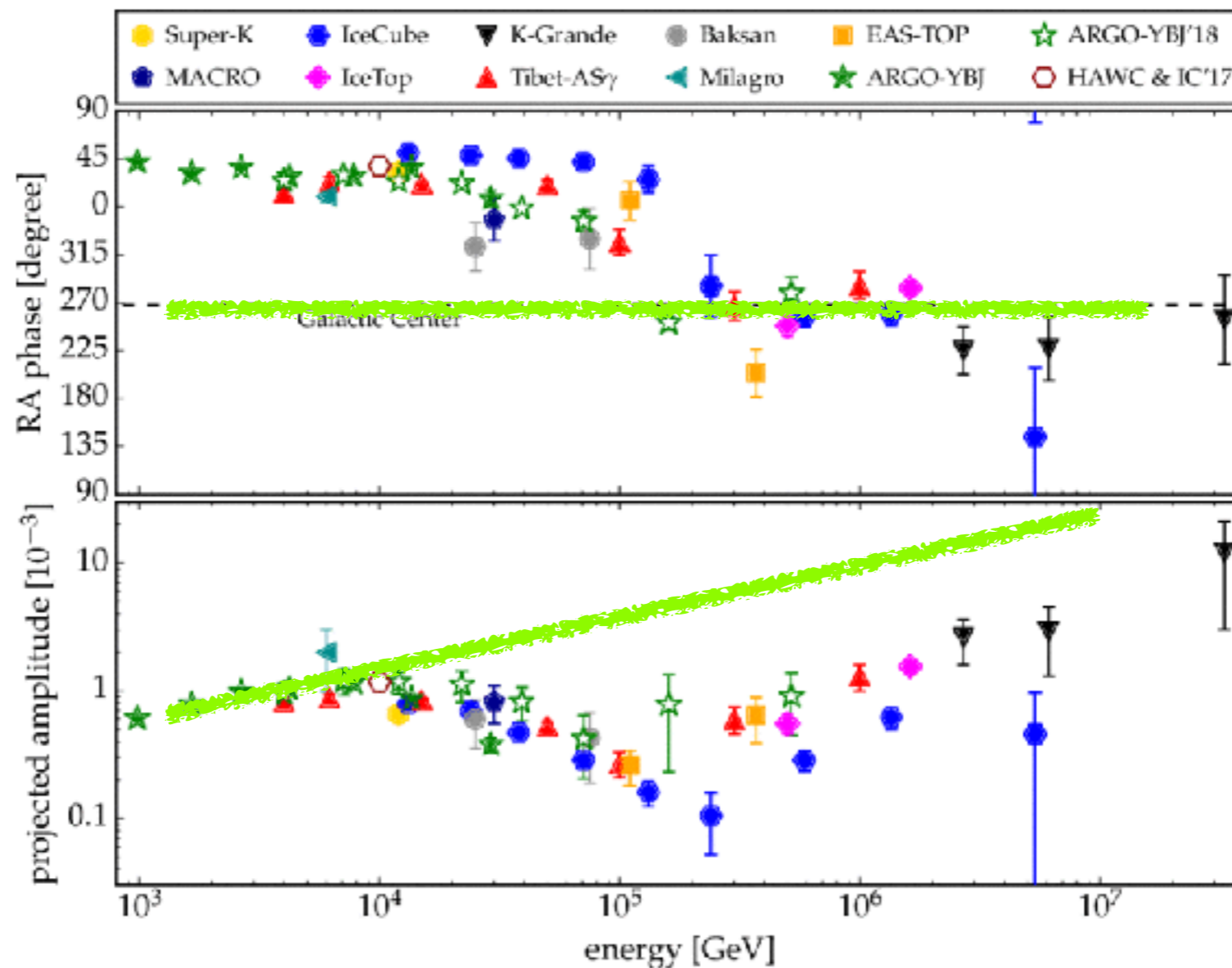
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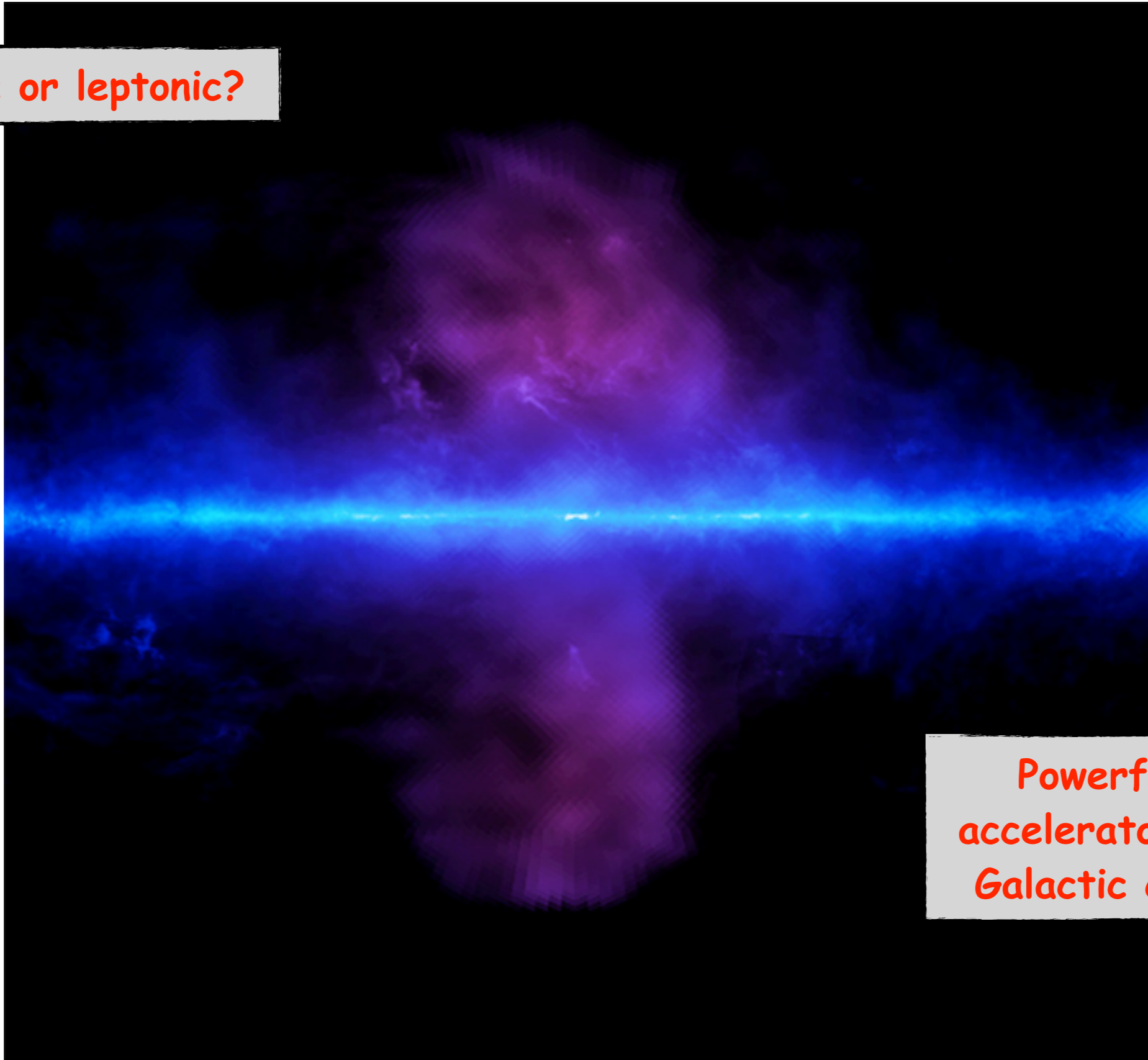
$$\int_{-1}^1 d\mu \mu v f_{\mu}^{(1)} = D \frac{\partial f^{(0)}}{\partial z}$$

points towards Galactic centre



Fermi bubbles

Hadronic or leptonic?



Powerful CR
accelerator in the
Galactic centre?

So? Two possibilities...

-> modify the two-power-laws model to account for all data (most popular)

-> invent a radically different scenario to account for all data (less popular)



Advertisements...

For a review of the problems of the SNR paradigm

Gabici et al. 2019, IJMPD, 28, 1930022-339 (arXiv:1903.11584)

PhD thesis on Cosmic Rays

COSMIC RAY ACCELERATION IN STAR FORMING REGIONS
(ED STEPUP)